

419(4) : The Three Kepler Laws in n Theory

Kepler's first law states by empirical observation that the orbit of m around M is an ellipse. In n theory this is replaced by:

$$\frac{dH}{dt} = 0 \quad - (1)$$

and

$$\frac{dL}{dt} = 0 \quad - (2)$$

giving many different types of orbit, precession, retrograde precession, shrinking and expanding orbits and so on.

Kepler's second law is derived by considering

Fig (1):

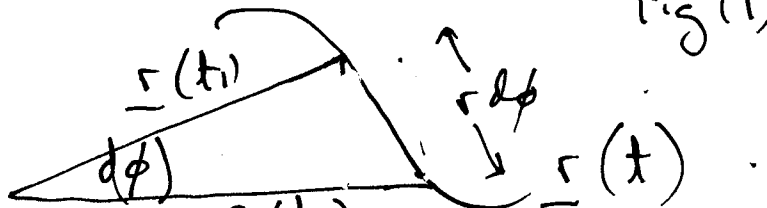


Fig (1)

so using the formula for the area of an isosceles triangle:

$$dA = \frac{1}{2} r^2 d\phi \quad - (3)$$

In Newtonian theory:

$$L = m r^2 \frac{d\phi}{dt} \quad - (4)$$

From eqs (3) and (4):

$$\frac{dA}{dt} = \frac{1}{2} r^2 \frac{d\phi}{dt} = \frac{L}{2m} = \text{constant} \quad - (5)$$

This is Kepler's second law of 1609, derived from data by Tycho Brahe.

In n theory, Kepler's second law is modified as follows. The angular momentum of n theory is:

$$L = \frac{\gamma m r \dot{\phi}}{m(r)} \quad - (6)$$

in which:

$$\gamma = \left(m(r) - \frac{v_N^2}{m(r)c^2} \right)^{-1/2} \quad - (7)$$

Therefore

$$\frac{d\phi}{dt} = \frac{m(r)L}{\gamma m r^2} \quad - (8)$$

so

$$\boxed{\frac{dA}{dt} = \frac{1}{2} \frac{m(r)L}{\gamma m}} \quad - (9)$$

It follows that:

$$\frac{dA}{dt} = \frac{L}{2m} m(r) \left(m(r) - \frac{v_N^2}{m(r)c^2} \right)^{1/2} \quad - (10)$$

where

$$v_N^2 \sim M G m(r)^{3/2} \left(\frac{2m(r)^{1/2}}{r} - \frac{1}{a} \right) \quad - (11)$$

In n theory, dA/dt is no longer constant.

Kepler's Third Law is the standard model
obtained from Eq. (5):

$$dt = \frac{2m}{L} dA \quad - (12)$$

so

$$T = \int_0^T dt = \frac{2m}{L} \int_0^A dA \quad - (13)$$

i.e.

$$T = \frac{2m}{L} A \quad - (14)$$

The time taken to complete one orbit, T , is
proportional to the area of the orbit, A . This is Kepler's
Third Law. In Newtonian theory A is the area of
an ellipse:

$$A = \pi ab \quad - (15)$$

where a and b are the semi major and semi minor axes:

$$a = \frac{d}{1-e^2}, \quad b = \frac{d}{(1-e^2)^{1/2}} \quad - (16)$$

where d is the half right latus rectum and e is the eccentricity.

so

$$T = \frac{2\pi m d^2}{L (1-e^2)^{3/2}} = \frac{2m\pi}{L} a^{3/2} d^{1/2} \quad - (17)$$

It follows that:

$$T^2 = 4\pi^2 \left(\frac{m^2 d}{L^2} \right) a^3 \quad - (18)$$

$$= \frac{4\pi^2}{mb} a^3$$

where we have used:

$$d = \frac{L^2}{m M b} \quad - (19)$$

Eq. (18) is Kepler's third law:

$$T^2 = \frac{4\pi^2}{mb} a^3 \quad - (20)$$

in the standard model.

In a theory Kepler's third law is modified as follows. From eq. (9):

$$dt = \frac{2m}{L} \frac{\gamma}{n(r)} dA \quad - (21)$$

so

$$T = \frac{2m}{L} \int_0^A \frac{\gamma}{n(r)} dA \quad - (22)$$

4) In the Newtonian limit:

$$\frac{\gamma}{n(r)} \rightarrow 1 \quad - (23)$$

and Kepler's third Law is recovered.

In Eq. (22):

$$\frac{\gamma}{n(r)} = \frac{1}{n(r)} \left(n(r) - \frac{v_N^2}{n(r)c^2} \right)^{-1/2} \quad - (24)$$

in which:

$$v_N^2 \sim \frac{M G m(r)}{r} \left(\frac{2m(r)}{r} - \frac{1}{a} \right) \quad - (25)$$

Therefore

$$\frac{\gamma}{n(r)} = f(r) \quad - (26)$$

i.e. $\gamma/n(r)$ is a function of r . In general,

assume that:

$$A = f_1(r), \quad - (27)$$

then:

$$dA = \frac{df_1(r)}{dr} dr \quad - (28)$$

and

$$T = \frac{2m}{L} \int_0^A \frac{\gamma}{n(r)} \frac{df_1(r)}{dr} dr \quad - (29)$$

$$T = \frac{2m}{L} \int_0^A f_1(r) \frac{\gamma}{n(r)} \frac{df_1(r)}{dr} dr \quad - (30)$$

If it is assumed that A is not a function of r , as for example is an elliptical orbit, where:

$$A = \pi ab \quad - (31)$$

then A is a constant of motion, independent of r .

Therefore assume that:

$$T \sim \frac{\gamma}{m(r)} \cdot \frac{2\pi A}{L} \quad - (32)$$

and

$$T^2 = \left(\frac{\gamma}{m(r)} \right)^2 \left(\frac{4\pi^2}{mG} a^3 \right) \quad - (33)$$

This is Kepler's third law in a theory.

From eqs. (24), (25) and (33) the time T can be evaluated in terms of $m(r)$ for a given M and a .

In the previous note it was found that the Newtonian limit

$$\frac{\gamma}{m(r)} \rightarrow 1 \quad - (34)$$

is wildly incorrect for the SS star. This means that the standard model mass for the "black hole" is completely incorrect. The effective mass in a theory:

$$M(\text{eff.}) = \left(\frac{m(r)}{\gamma} \right)^2 M_N \quad - (35)$$

where M_N is the Newtonian mass:

$$M_N = \frac{4\pi^2 a^3}{T^2 G} \quad - (36)$$